

Two very different mhd codes for astrophysics: Athena and GIZMO.



Scope and limitations



Basic equations to solve

Equation of mass continuity $\frac{\partial \rho}{\partial t} + \nabla \cdot (\rho \mathbf{v}) = 0$

Equation of motion $\frac{\partial \rho \mathbf{v}}{\partial t} + \nabla \cdot (\rho \mathbf{v} \mathbf{v} - \mathbf{B} \mathbf{B} + P) = 0$

Equation of energy $\frac{\partial E}{\partial t} + \nabla \cdot ((E + P) \mathbf{v} - \mathbf{B} (\mathbf{B} \cdot \mathbf{v})) = 0$

Induction equation $\frac{\partial \mathbf{B}}{\partial t} - \nabla \times (\mathbf{v} \times \mathbf{B}) = 0$

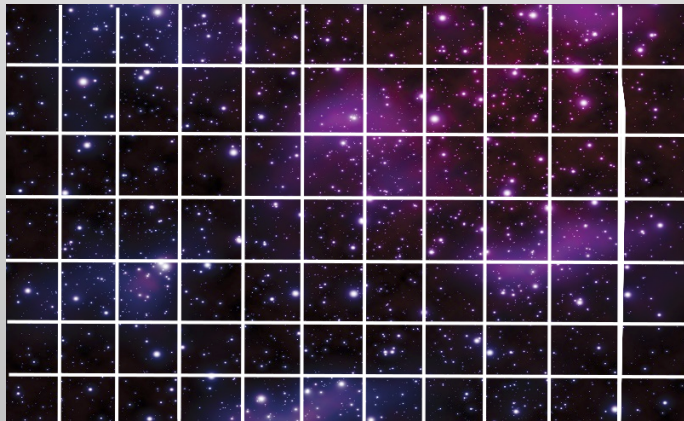
$$P = P_{hydro} + \frac{\mathbf{B} \cdot \mathbf{B}}{2}$$

$$E = \frac{P_{hydro}}{\gamma - 1} + \frac{\rho (\mathbf{v} \cdot \mathbf{v})}{2} + \frac{\mathbf{B} \cdot \mathbf{B}}{2}$$

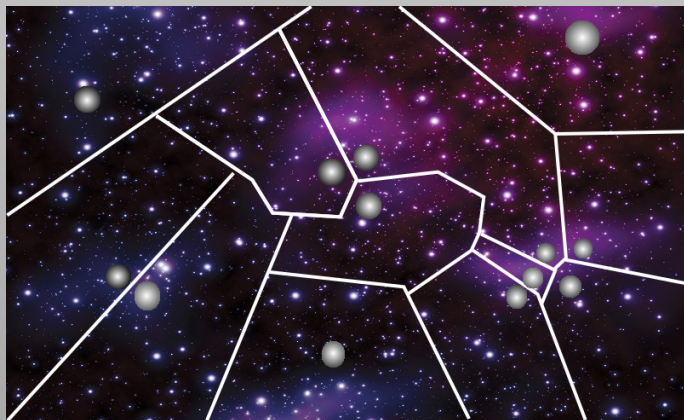
(No microscopic dissipation)



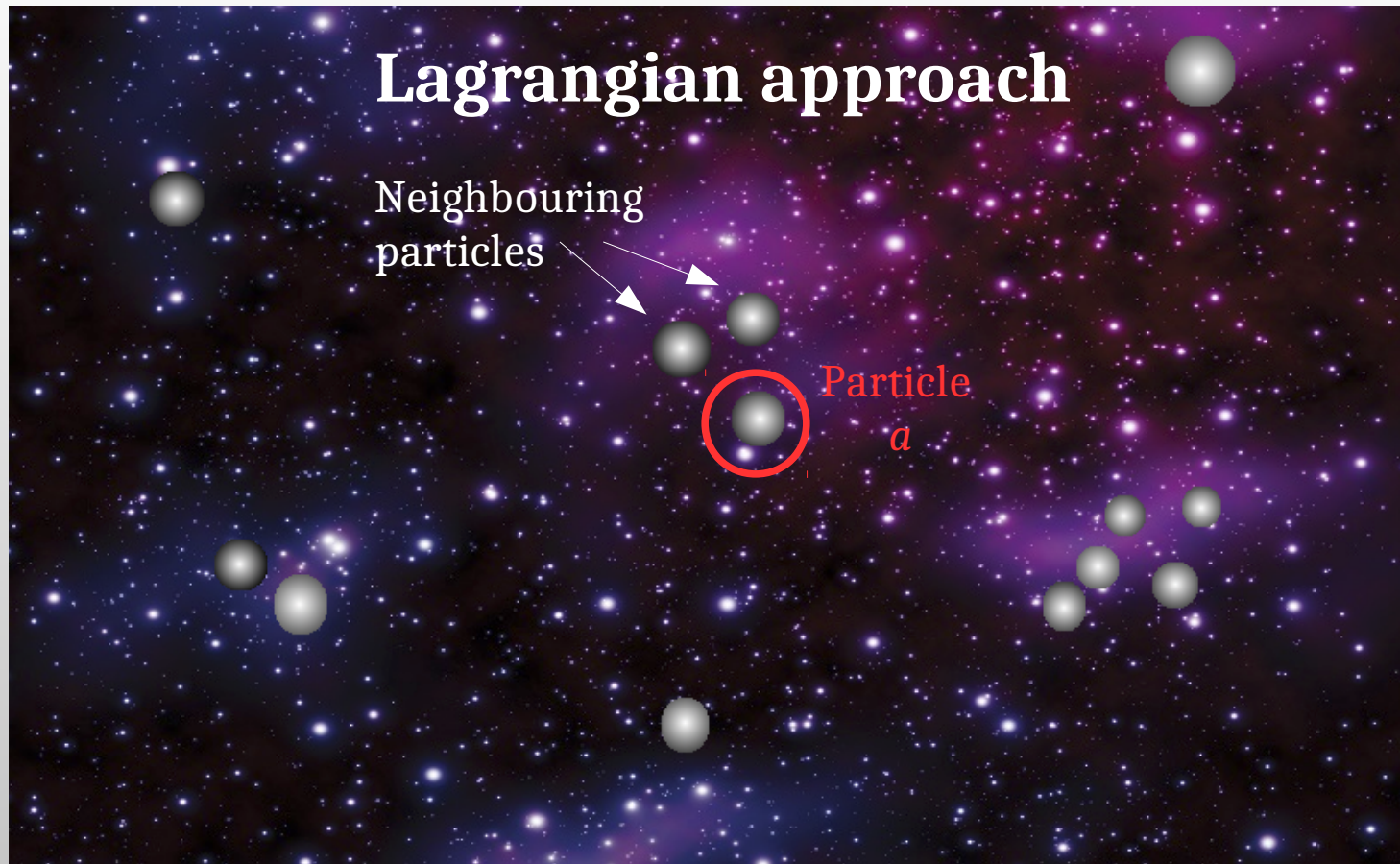
Smooth Particle
(magneto)Hydrodynamics
SPH, SPMHD



Grid-based codes
Finite differences
Godunov methods
(finite volumes)



Moving
mesh/mesh-free +
Godunov methods

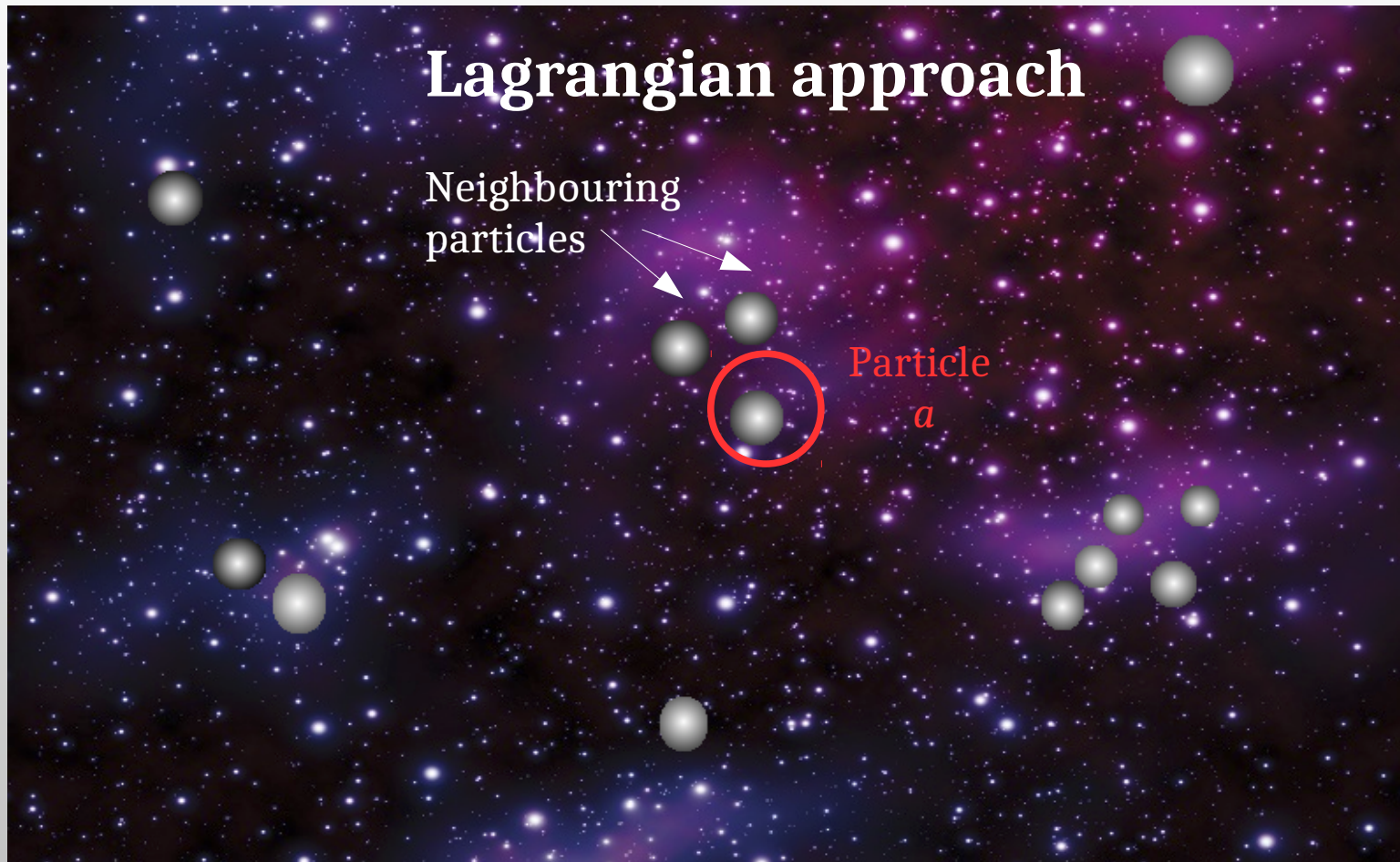


Since 1977 (*Lucy, 1977; Gringold & Monaghan, 1977*)

Fluid seen as a set of particles of mass m that are moved with the local fluid velocity $\mathbf{v}$

$$\rho_a = \sum_b m_b W(|\mathbf{r}_a - \mathbf{r}_b|, h_a)$$

Sum over neighbouring particles inside $R_{\text{kern}} h$



PHANTOM (Price et al. 2017)

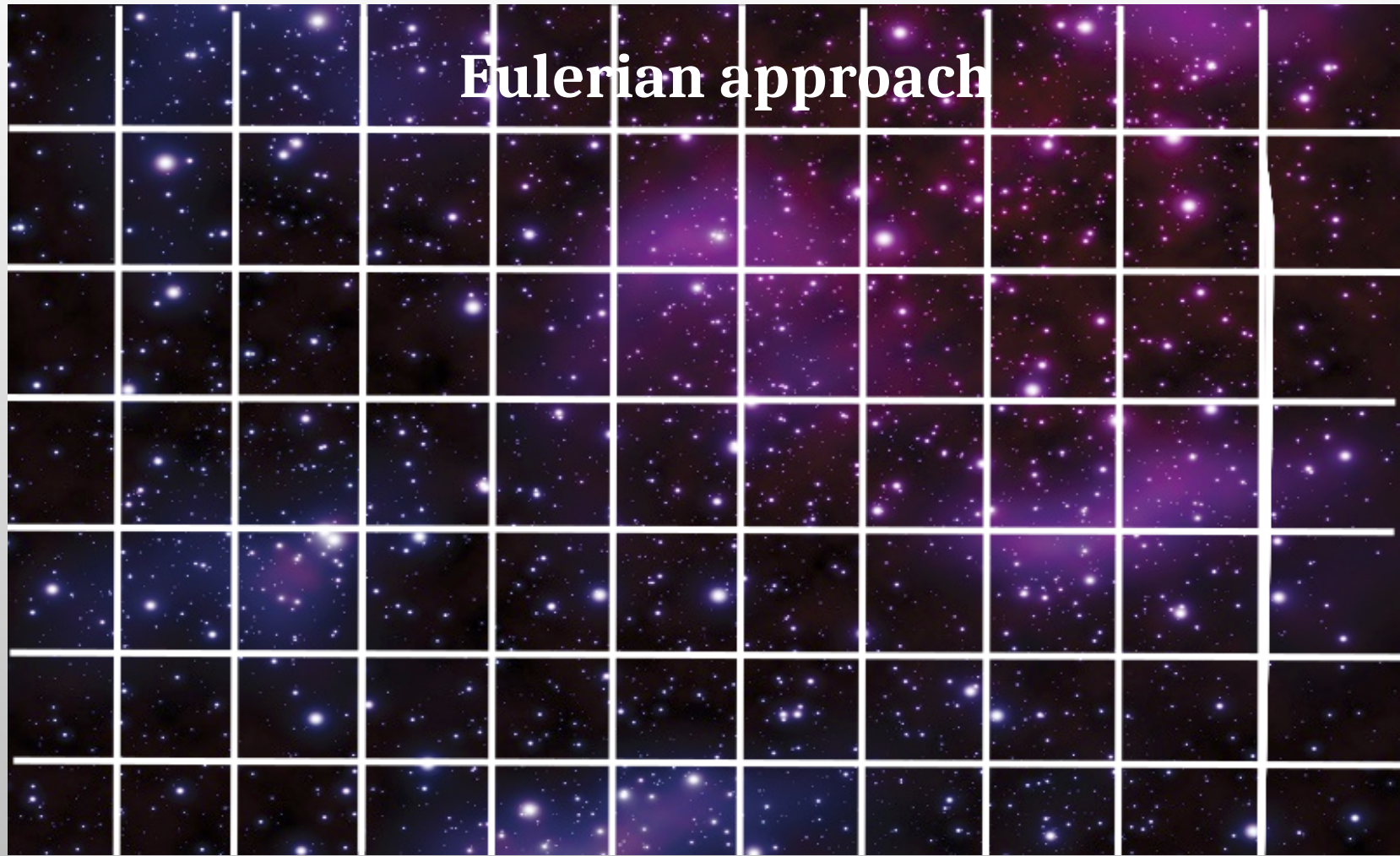
SPMHD code

<https://phantomsph.bitbucket.io/>

Last stable release: 13/03/2018

Developed over the last decade, Fortran 90





Volume discretised into cells

Fluid equations solved across the area/volume elements

Finite-volume Godunov schemes

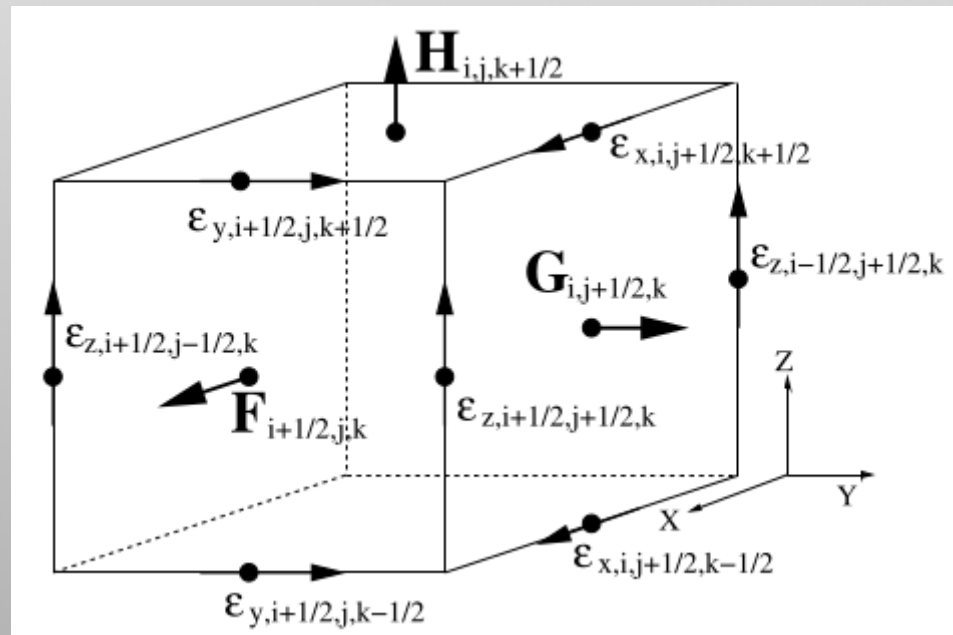
ATHENA (*Stone et al., 2008*)

C language

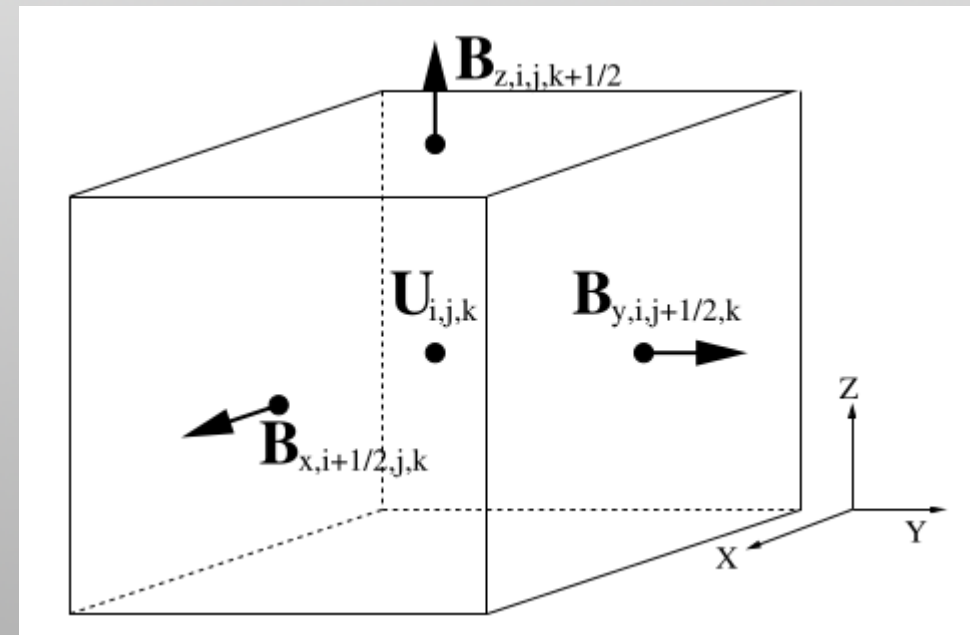
Publicly available. Latest version (v4.2) from 2013

<https://princetonuniversity.github.io/Athena-Cversion/AthenaDocsDownLd>

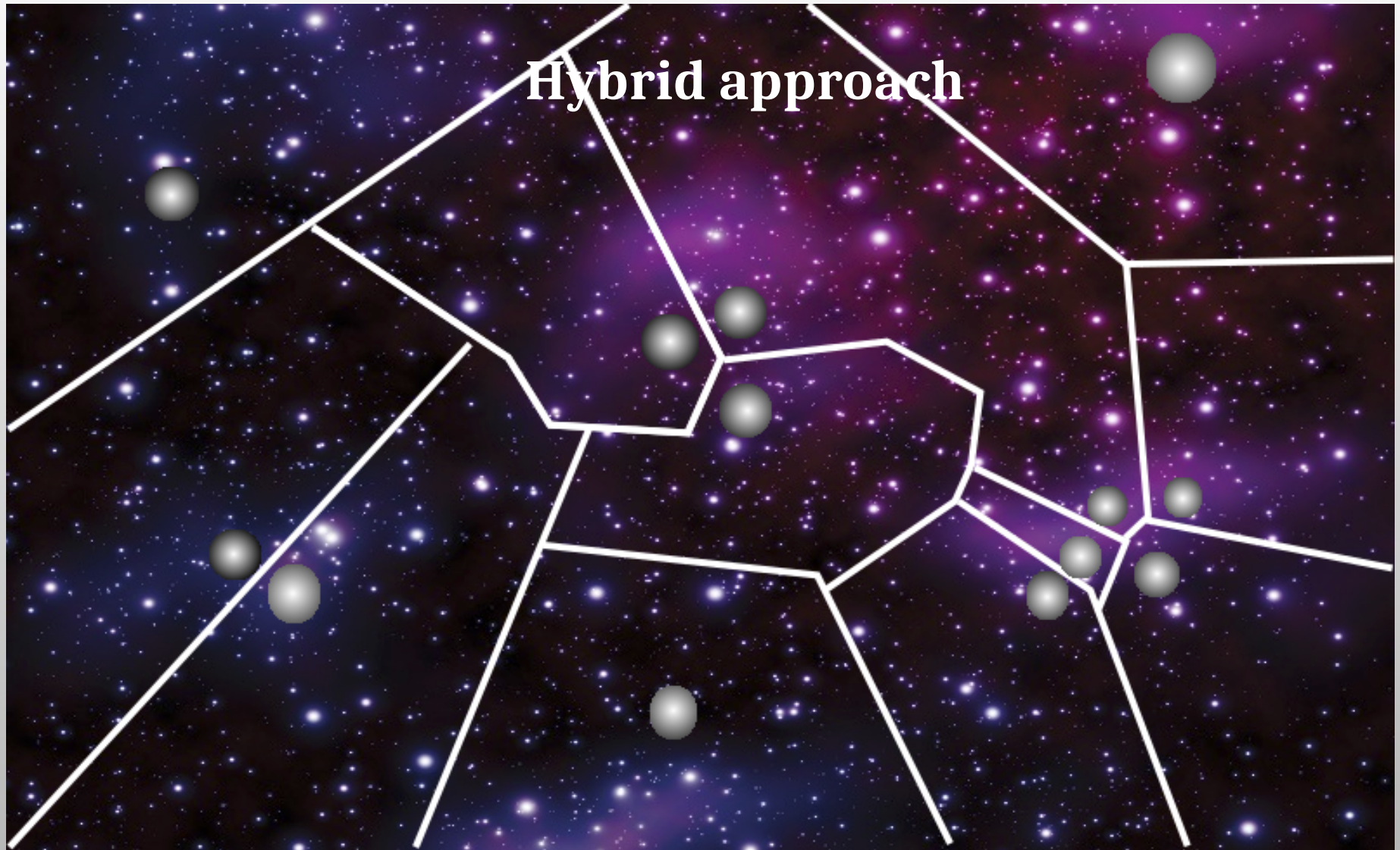
Documentation <https://princetonuniversity.github.io/Athena-Cversion/>



Mass, momentum and energy fluxes:
Finite volumes



Magnetic fields:
finite areas



Non-regular mesh

Cells move and deform continuously

Lagrangian + Finite-volume Godunov methods

GIZMO *(Hopkins, 2015)*

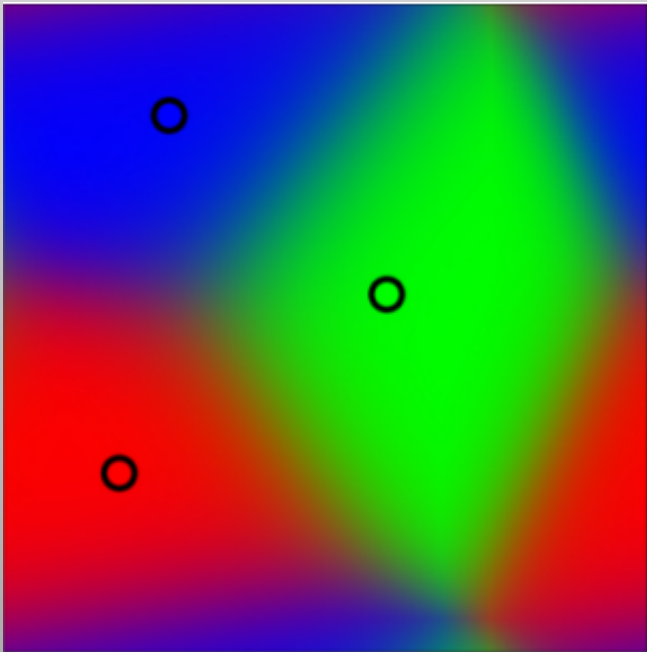
C language

Publicly available. Latest version 18/02/2018

<https://bitbucket.org/phopkins/gizmo-public>

Documentation:

http://www.tapir.caltech.edu/~phopkins/Site/GIZMO_files/gizmo_documentation.html



Lagrangian/Eulerian code
Meshless Finite Volume (MFV) and
Meshless Finite Mass (MFM) methods

More about Athena



ATHENA

Implemented physics:

- Compressible HD/MHD in 1D/2D/3D
- Special Relativity HD and MHD
- Ideal gas equation of state (adiabatic and isothermal)
- Passive scalars advected with the flow
- Self gravity, static gravitational potential
- Ohmic resistivity, ambipolar diffusion and Hall effect
- Navier-Stokes and anisotropic (Braginskii) viscosity
- Isotropic/anisotropic thermal conduction
- Optically-thin radiative cooling
- Aerodynamic particles

Additional features:

- Cartesian/cylindric coordinates
- Static Mesh Refinement
- Shearing Box and orbital advection algorithm for MHD
- Parallelization

2004 – 2008: Private use

Magnetothermal Instability (MTI), Magnetorotational Instability (MTI),
Protoplanetary Discs (PPDs), Molecular Clouds (MC), MHD turbulence.

Algorithms development (*Gardiner & Stone, 2005, 2008*)

2008: Public version release (*Stone et al. 2008*)

Heat conduction (*Parrish, Stone and Lemaster, 2008*)

2009

Viscosity (*Dong & Stone, 2009*)

Resistivity (*Fromang & Stone, 2009*)

Charged particles (*Lehe, Parrish, Quataert, 2009*)

2010

Cylindrical coordinates (*Skinner & Ostriker, 2010*)

Shearing box and orbital advection scheme (*Stone & Gardiner, 2010*)

Aerodynamic particles (*Bai & Stone, 2010*)

2011

Self-gravity (*Gong & Ostriker, 2011*)

Relativistic MHD (*Beckwith & Stone, 2011*)

Ambipolar diffusion (*Bai & Stone, 2011*)

2012

Radiative MHD (*Gendeleev & Krumholtz, 2012*)

Radiative cooling, heating and heat conduction (*Choi & Stone, 2012*)

Radiative transfer (*Davis, Stone & Jiang, 2012*)

Orbital advection in cylindrical coordinates (*Sorathia et al., 2012*)

Anisotropic conduction and Braginskii viscosity (*Kunz et al., 2012*)

2013

Sink particles (*Gong & Ostriker, 2013*)

Hyperion: Two-moment Radiation HD (*Skinner & Ostriker, 2013*)

2014

Pegasus: hybrid-kinetic particle-in-cell code (Kunz, Stone & Bai, 2014)

Dust particles in cylindrical coordinates (Zhu et al., 2014)

Another algorithm for radiation MHD (Jiang, Stone & Davis, 2014)

2017

TIGRESS, Three phase Interstellar Medium in Galaxies Resolving Evolution with Star formation and Supernova Feedback (Kim & Ostriker, 2017)

Instabilities

Turbulence

Protoplanetary
&
Accretion Discs

Interstellar
Medium

Extragalactic

Radiative
HD/MHD

Plasmas

Relativistic
HD/MHD

- Magnetorotational Instability
- Ambipolar diffusion
- Resistivity
- Hall effect
- Turbulence in magnetised discs
- Diffusion/dynamic of solids
- Concentration of particles in vortices/dusty vortices
- Planet-disc interactions
- Angular momentum transport
- Winds

- Super Novae feedback (radiation, chemistry)
- Molecular cloud formation
- Superbubbles
- Shocks
- Prestellar/magnetised cores
- Dusty HII regions
- Striations
- Particular regions: Virgo cluster, Vela C
- Dust polarization

- Galactic clusters: magnetic fields, turbulence
- Spiral shocks
- Barred galaxies
- Turbulence in discs
- Ram pressure stripping
- AGNs: jets, clouds, accretion disc, magnetic fields
- Galactic superbubbles
- Galactic winds
- Disc-corona interface
- Circumgalactic medium and feedback
- Formation of black holes
- Star formation

- Magnetorotational
- Thermal/magnetothermal
- Streaming Inst. (*Youdin & Goodman, 2005*)
- Current-driven
- Heat-flux Buoyancy inst. (*Quataert, 2008*)
- Wiggle Instability (*Wada & Koda, 2004*)
- Pulsating Instability of fast turbulent flames
- Kruskal-Schwarzschild
- Zombie Vortex (dusty vortex) (*Marcus et al., 2015*)

- MHD
- Turbulent flames
- Alfvénic turbulence
- Weak turbulence
- Galactic clusters
- Supersonic isothermal turbulence

- Radiation feedback in ULIRGS
- Hot accretion disc coronae
- Photoevaporative mass loss from hot jupiters
- Massive star envelopes
- Iron opacity bump
- Chemistry and radiative shielding in star forming galactic discs
- Dust-driven winds
- Dusty cloud acceleration

- Magnetized Noh pinch problem
- HII regions
- Buoyancy instabilities
- Contact instabilities
- Weakly compressible MHD turbulence
- Coupling cosmic rays and thermal plasma
- Particle acceleration by a solar flare termination shock

- Magnetised jets
- Variability in AGNs
- Acceleration of relativistic electrons
- Kruskal-Schwarzschild instability

More about GIZMO



GIZMO

Implemented physics:

- Compressible HD/MHD in 1D/2D/3D
- Cosmological integrations
- Radiative heating/cooling and chemistry
- Galaxy/Star/Black hole formation and feedback
- Non-standard cosmology
- Self-gravity with adaptive gravitational resolution
- Anisotropic conduction and viscosity
- Passive scalars, turbulent eddy diffusion, shearing box
- **Dusty fluids**
- Degenerate/stellar equations of state
- Elastic/plastic/solid-body dynamics
- Radiation hydrodynamics

Science: mainly cosmology, AGNs, galaxies

MHD implementation *(Hopkins & Raives, 2016)*

Solution for the divergence problem: Constrained Transport (CT) method *(Evans & Hawley, 1988)*

CT difficult to implement in Lagrangian codes

How do they solve it: Powell 8-wave cleaning + Dedner cleaning

Powell 8-wave: subtracts unstable error terms resulting from non-zero divergence from the eq. of motion

Dedner cleaning: adds source terms that transport the divergence away and then damp it.

$$\log\left(\frac{h_i |\nabla \cdot \mathbf{B}_i|}{|\mathbf{B}_i|}\right) \in [-7, -2], \sim -4 \quad h_i = \text{effective cell size}$$

Charged grains *(Lee et al., 2017)*

$$\frac{d\mathbf{u}_d}{dt} = \frac{-\mathbf{u}_d - \mathbf{u}_{gas}}{t_s} + \frac{Z_d e}{m_d c} (\mathbf{u}_d - \mathbf{u}_{gas}) \times \mathbf{B}$$

Dust-dust collisions not included

Study the interaction of Alfvén waves with Molecular Clouds

Five fluids

neutrals, ions, electrons, neutral grains, charged grains

Effects

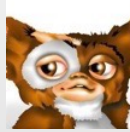
ambipolar diffusion, interaction of charged species with charged grains

Athena or GIZMO?



Key factors to be considered

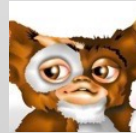
Popularity



Athena (407 citations in 10 years), GIZMO (177 citations in 3 years)

Efficiency/computational cost

Available modules



Support



Google group

Time



Spent learning Athena + programming a module vs learning a new method

$\nabla \cdot \mathbf{B} = 0$



Particle collisions $\emptyset$

- Bai & Stone (2010), ApJS, 190:297
Bai & Stone, (2011) ApJ, 736:144
Beckwith & Stone (2011) ApJS, 193:6
Choi & Stone (2012) ApJ, 747:86
Davis, Stone & Jiang (2012) ApJS, 199:9
Dong & Stone (2009) ApJ 704:1309
Evans & Hawley (1988) ApJ 332:659
Fromang & Stone (2009) A&A, 507:19
Gardiner & Stone (2005) JCoPh, 205:509
Gardiner & Stone (2008) JcoPh, 227:4123
Gendeleev & Krumholtz (2012) ApJ, 745:158
Gong & Ostriker (2011) ApJ, 190:297
Gong & Ostriker (2013) ApJS, 204:8
Gingold & Monaghan (1977) MNRAS, 181:375
Hopkins (2015) MNRAS, 450:53
Hopkins & Raives (2016) MNRAS, 455:51
Jiang, Stone & Davis (2014) ApJS, 213:7
Kim & Ostriker (2017) ApJ, 846:133

- Kunz, Bogdanović, Reynolds et al. (2012) *ApJ*, 754:122
- Kunz, Stone & Bai (2014) *JcoPh*, 259:154
- Lee, Hopkins & Squire (2017) *MNRAS*, 469:3532
- Lehe, Parrish, Quataert (2009) *ApJ*, 707:404
- Lucy (1977) *AJ*, 82:1013
- Marcus, Pei, Jiang et al. (2015) *ApJ*, 808:87
- Parrish, Stone and Lemaster (2008) *ApJ*, 688:905
- Price, Wurster, Nixon et al. (2017) *arXiv*: 1702.03930v1
- Quataert (2008) *ApJ*, 673:758
- Skinner & Ostriker (2010) *ApJS*, 188:290
- Skinner & Ostriker (2013) *ApJS*, 206:21
- Sorathia, Reynolds, Stone et al. (2012) *ApJ*, 749:189
- Stone, Gardiner, Teuben et al. (2008) *ApJS*, 178:137
- Stone & Gardiner (2010) *ApJS*, 189:142
- Wada & Koda (2004) *MNRAS*, 349:270
- Youdin & Goodman (2005), *ApJ*, 620:459
- Zhu, Stone, Rafikov et al. (2014) *ApJ*, 785:122